Screening mechanism and dark coupling

Kazuya Koyama



Institute of Cosmology and Gravitation University of Portsmouth

General relativity

Why do we believe general relativity?

Observational point of view

GR is tested to very high accuracy by solar system experiments and pulsar timing measurements

C.Will gr-qc/0510072

Theoretical point of view

GR is the unique metric theory in 4D that gives second order differential equations





FIFTH FORCE ENERGY IS THE BUILDING BLOCK ON WHICH ALL LIFE IS CREATED AND IS THE ESSENCE OF LIFE ITSELF. FIFTH FORCE IS IN FACT MATTER BUT IT CANNOT BE SEEN AND THIS MATTER IS THE NATURE OF THE PRIME. OUT OF ALL THE ENERGIES IN THE UNIVERSE FIFTH FORCE IS THE MOST POWERFUL ENERGY OF ALL BECAUSE IT IS IN EVERY LIFE THING AND IT IS IN EVERY LIFE BEING WHETHER THEY ARE OF THE HUMAN RACE OR AN ALIEN RACE LIVING ON A PLANET IN A FAR AND DISTANT GALAXY IN OUR UNIVERSE OR IN A PARALLEL UNIVERSE LIGHT YEARS WAY.

Brans-Dicke theory

Action

$$S = \int d^4 x \left(\psi R - \frac{\omega_{BD}}{\psi} \left(\nabla \psi \right)^2 + V(\psi) \right) \qquad V \sim H_0^2 M_{pl}^2$$

$$f(R)$$
 gravity: $\omega_{BD} = 0$

quasi-static approximations (neglecting time derivatives) $ds^{2} = -(1+2\Psi)dt^{2} + a(t)^{2}(1-2\Phi)d\bar{x}^{2} \quad \Psi = \Psi_{0} + \varphi$ $(3+2\omega_{BD})\nabla^{2}\varphi = -8\pi G\rho$ $\nabla^{2}\Psi = 4\pi G\rho - \frac{1}{2}\nabla^{2}\varphi$ $\Phi - \Psi = -\varphi$

Constraints on BD parameter

Solutions

$$(3+2\omega_{BD})\nabla^{2}\varphi = -8\pi G\rho$$

$$\nabla^{2}\Psi = -4\pi G \left(\frac{4+2\omega_{BD}}{3+2\omega_{BD}}\right)\rho, \quad G_{eff} = \left(\frac{4+2\omega_{BD}}{3+2\omega_{BD}}\right)G$$

$$\Psi = \frac{2+\omega_{BD}}{1+\omega_{BD}}\Phi \equiv \gamma^{-1}\Phi$$

PPN parameter

$$\gamma = \frac{1 + \omega_{BD}}{2 + \omega_{BD}}$$

$$\gamma - 1 = (2.1 \pm 2.3) \times 10^{-5} \qquad \omega_{BD} \ge 40,000$$

This constraint excludes any detectable modifications in cosmology

How we suppress fifth force (1)

Break equivalent principle
Einstein frame $\begin{aligned}
\alpha &= \sqrt{\frac{1}{3+2\omega_{BD}}} \\
S_E &= \int d^4 x \left[\sqrt{-\overline{g}} \left(\overline{R} - \frac{1}{2} (\overline{\nabla} \phi)^2 + \overline{V}(\phi) \right) + L_m [e^{-\alpha \phi/M_{pl}} \overline{g}_{\mu\nu}] \right] \\
\text{Wetterich , Amendora Baldi's talk} \\
S_E &= \int d^4 x \left[\sqrt{-\overline{g}} \left(\overline{R} - \frac{1}{2} (\overline{\nabla} \phi)^2 + \overline{V}(\phi) \right) + L_{DM} [e^{-\alpha \phi/M_{pl}} \overline{g}_{\mu\nu}] + L_{baryon} [\overline{g}_{\mu\nu}] \right] \\
\text{Jordan frame}
\end{aligned}$

$$S = \int d^4x \left(\psi R - \frac{\omega_{BD}}{\psi} \left(\nabla \psi \right)^2 + V(\psi) + L_{DM} [g_{\mu\nu}] + L_{baryon} [f(\psi)g_{\mu\nu}] \right)$$

the coupling compensates the fifth force acting on baryons (however, interpretation of observations is non-trivial)

Deruelle, Sasaki 1012.5386

How to suppress the fifth force (2)

$$S = \int d^4x \left(\psi R - \frac{\omega_{BD}(\psi)}{\psi} \left(\nabla \psi \right)^2 + V(\psi) + L_m[g] \right)$$

GR is recovered if (i) the mass is large $V'' \rightarrow \infty$ or (ii) the kinetic term is large $\omega_{BD} \rightarrow \infty$

These limits should be realised in environmentally (density) dependent way to avoid the recovery of GR on all scales

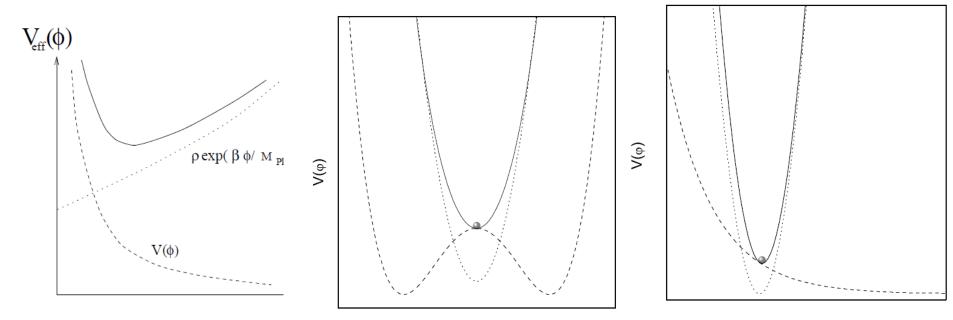


Chameleon/symmetron/dilaton

Talks by Brax, Davis, Khoury, Li, Hu,..

$$S_{E} = \int d^{4}x \left[\sqrt{-\overline{g}} \left(\overline{R} - \frac{1}{2} \left(\overline{\nabla} \phi \right)^{2} + \overline{V}(\phi) \right) + L_{m} [A(\phi)^{2} \overline{g}_{\mu\nu}] \right]$$
$$\Box \phi = V_{eff}(\phi), \quad V_{eff}(\phi) = \overline{V}(\phi) - (A(\phi) - 1)\rho$$

Einstein frame



Local tests forbid the modification[®] above a few Mpc Khour[®]y's talk

How to suppress the fifth force (3)

• Vainshtein mechanism originally discussed in massive gravity rediscovered in DGP brane world model linear theory $\omega_{BD} = 0$ $3\nabla^2 \varphi = -8\pi G \rho$ $\nabla^2 \Psi = 4\pi G \rho - \frac{1}{2} \nabla^2 \varphi$

even if gravity is weak, the scalar can be non-linear Talks by Hu, Hui $3\nabla^2 \varphi + r_c^2 \left\{ \left(\nabla^2 \varphi \right)^2 - \partial_i \partial_j \varphi \, \partial^i \partial^j \varphi \right\} = 8\pi G a^2 \rho \qquad r_c \sim m^{-1} \sim H_0^{-1}$

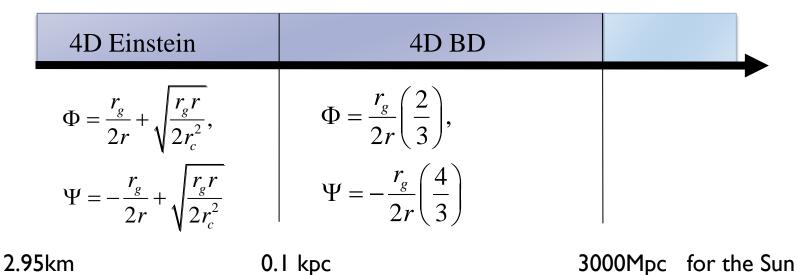
Vainshtein mechanism

Spherically symmetric solution for the scalar

$$\frac{d\varphi}{dr} = \frac{r_g}{r^2} \Delta(r), \quad \Delta(r) = \frac{2}{3} \left(\frac{r}{r_V}\right)^3 \left(\sqrt{1 + \left(\frac{r_V}{r}\right)^3} - 1\right) \quad r_V = \left(\frac{8r_c^2 r_g}{9}\right)^{\frac{1}{3}}, \quad r_g = 2GM$$

1





Solar system constraints

• The fractional change in the gravitational potential $\varepsilon = \frac{\partial \Psi}{\Psi}$ The anomalous perihelion precession

$$\delta\phi = \pi r \frac{\partial}{\partial r} \left[r^2 \frac{\partial}{\partial r} \left(\frac{\varepsilon}{r} \right) \right]$$

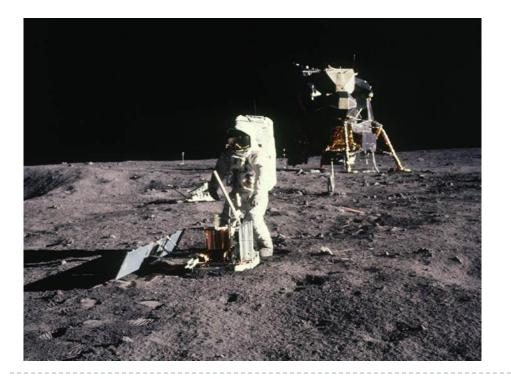
The vainshtein radius is shorter for a smaller object Lunar laser ranging: the Erath-moon distance $r_{E-M} = 4.1 \times 10^5 \text{ km}$

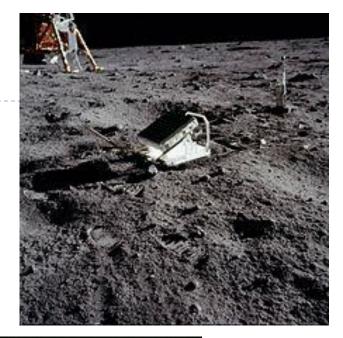
$$\delta\phi = \frac{3\pi M_{pl}}{4} \left(\frac{r_{E-M}^{3}}{8\pi r_{c}^{2} M_{\oplus}}\right)^{1/2} < 2.4 \times 10^{-11}, \quad r_{E-M} << r_{V}$$

$$r_c > 10^{60} M_{pl}^{-1} \sim H_0^{-1}$$
 Dvali et.al 0212069
Burrgae & Seery 1005.1927

Local tests are beginning to exclude the acceleration from new physics









Observational implications (1) equivalence violation

(1) Equivalence violation
 interacting dark energy model (Einstein frame)
 modified gravity models (Jordon frame)

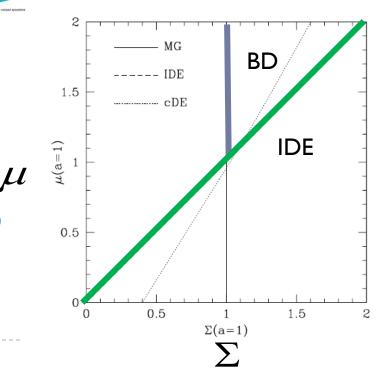
$$k^{2}(\Psi + \Phi) = -4\pi Ga^{2}\Sigma(a,k)\rho_{m}\delta_{m}$$

$$k^{2}\Psi = -4\pi Ga^{2}\mu(a,k)\rho_{m}\delta_{m}$$
BD
IDE
$$\Sigma(a,k) \approx 1$$

$$\Sigma = \mu$$

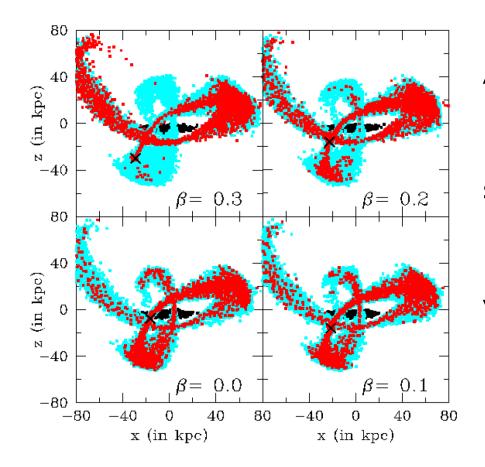
$$\mu(a,k) \approx \frac{2(2+\omega_{BD})}{3+2\omega_{BD}}$$
($\Phi = \Psi$)

A combination of weak lensing and RSD can break the degeneracy (i.e. Eg parameter) Reyes's talk



Observational implications (1) equivalence violation

Tidal disruption of satellite galaxies



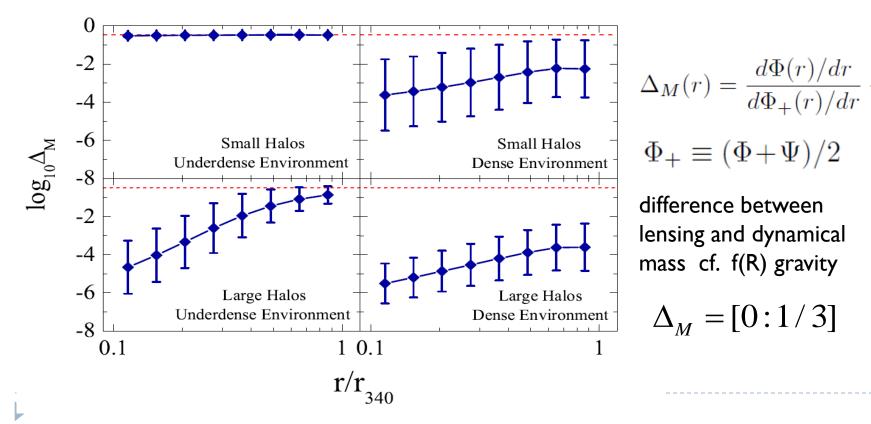
Kesden, Kamionkowski astro-ph/0608095

An enhanced force acting on DM redistribute stars from the leading stream to the trailing stream

10% enhancement of DM force would be excluded

Observational implication (2) Environmental screening

 (2) Environmental dependent screening Zhao's talk dark matter halos
 Schmidt 1003.0409, Zhao, Li Koyama 1011.1257 Li, Zhao, Koyama 1111.2602
 screening depends on environment and halo mass

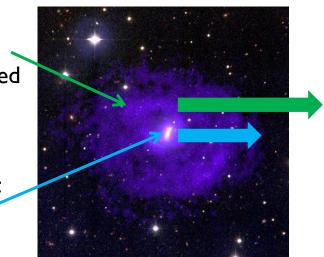


Observational implication (2) Environmental screening

Apparent violation of equivalent principle

HI gas:

Stellar disk: screened



Hui, Nicolis, Stubbs 0905.2966 Jain & Vander Plas 1106.0065



• The rotation curve of HI gas is enhanced compare with stars

• The stellar disk is displaced from the HI gas disk

If we can use dwarf galaxies in voids that have shallower potentials than the Milky way, we can go beyond local tests

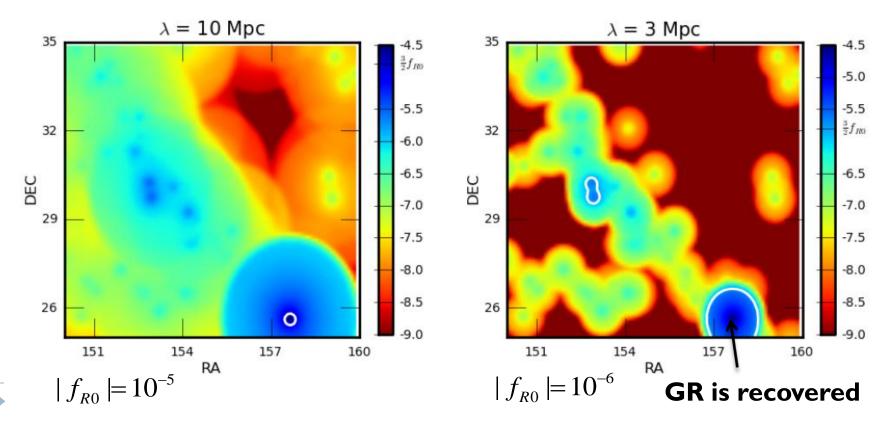
Creating a screening map

It is essential to find places where GR is not recovered

Small galaxies in underdense regions Cabre, Vikram, Zhao, Jain, KK

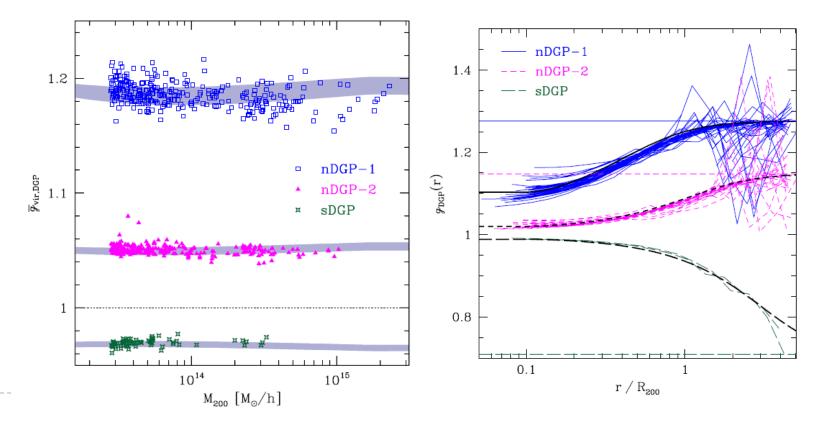
1204.6046

SDSS galaxies within 200 Mpc



- (3) Vainshtein mechanism
 - dark matter halos Schmidt 1003.0409

Screening depends on environment and mass very weakly



Morphology dependence
 The non-linear term vanishes for ID plane wave

 $\left(\nabla^2\varphi\right)^2 - \partial_i\partial_j\varphi\,\partial^i\partial^j\varphi$

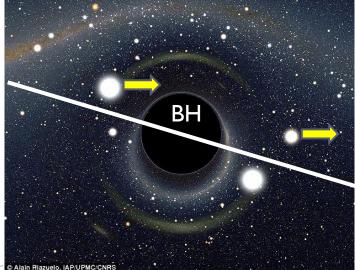
screening is weak in filaments



Apparent equivalent principle violation Hui, Nicolis 1201.1508

stars can feel an external field generated by large scale structure but a black hole does not due to no hair theorem Hui's talk

central BH lag behind stars



Non-superposition Hiramatsu, Hu, KK, Schmidt

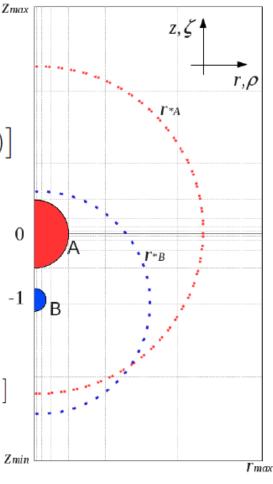
$$\nabla^2 \phi = g(a)a^2 \left(8\pi G\delta\rho - N[\phi,\phi]\right)$$

$$N[\phi_A, \phi_B] = \frac{r_c^2}{a^4} \left[(\nabla^2 \phi_A \nabla^2 \phi_B - (\nabla_i \nabla_j \phi_A) (\nabla^i \nabla^j \phi_B) \right]$$

Two body problem (cf. Earth-moon)

$$\phi = \phi_A + \phi_B + \phi_\Delta$$

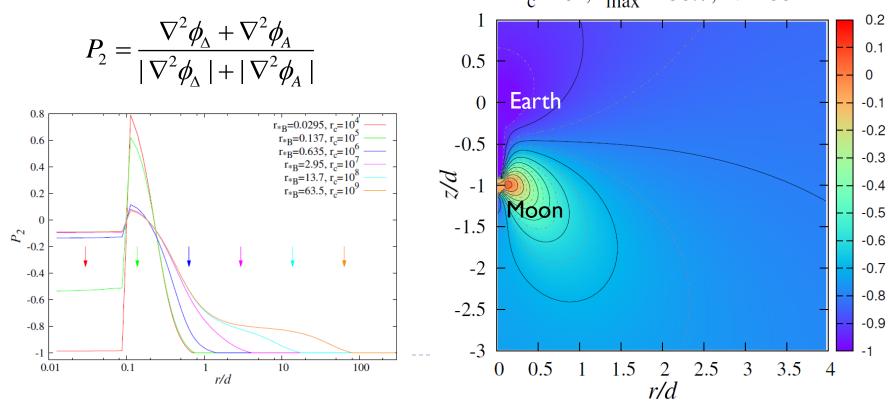
$$\nabla^2 \phi_\Delta + N[\phi_\Delta, \phi_\Delta] + 2N[\phi_A + \phi_B, \phi_\Delta] = -2N[\phi_A, \phi_B]$$



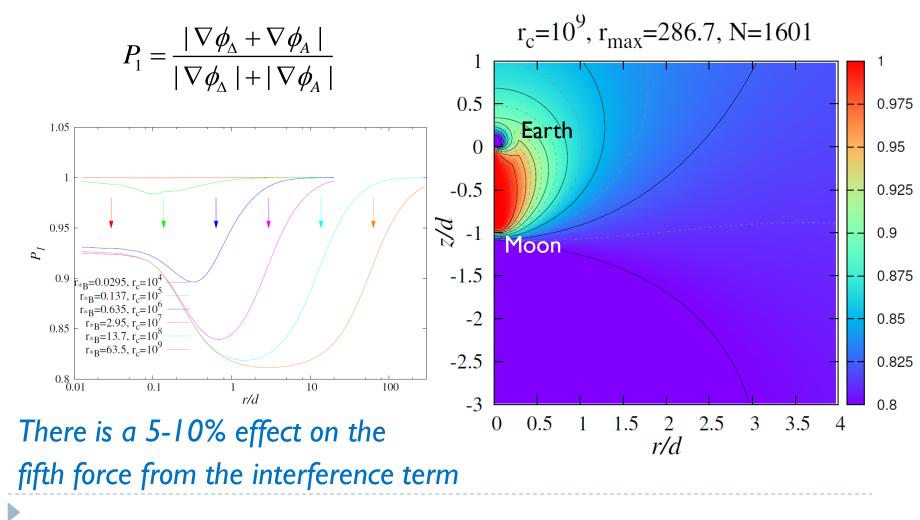
Near a small body B (moon)

$$\nabla^2 \phi_{\Delta} = -\nabla^2 \phi_A + O(\sqrt{M_A / M_B})$$

the interference term cancels the second derivative of the field from the large body (Earth) $r_c=10^9$, $r_{max}=286.7$, N=1601



Observational implication (3) Vainshtein mechanism
Effects on the first derivative (force) is small



Challenge for simulations

 Screening mechanism governed by a non-linear Poisson equation

 $\nabla^2 \phi = 4\pi G A(\phi) \rho + V'(\phi) + N[\partial \phi, \partial^2 \phi]$

no superposition rule it is not possible to separate long and short range forces need to solve the non-linear Poisson equation on a mesh

MLAPM

Li, Zhao 0906.3880, Li, Barrow 1005.4231 Zhao, Li, Koyama 1011.1257

ECOSMOG

(based on
RAMSES)Li, Zhao, Teyssier, Koyama 1110.1379
Jennings et.al. 1205.2698, Li et.al. 12064317
Brax et.al. 1206.3568

ECOSMOG-Vainshtein simulations are underway Li, Zhao, Koyama

Schimidt 0905.0858, 0910.0235 Chan & Scoccimarro 0906.4548

Conclusion

- Modification of GR generally introduce the fifth force, which should be screened
 - I) break equivalence principle and remove coupling to baryons
 Einstein frame interacting dark energy models
 - 2) Environmentally (density) dependent screening Chameleon/Symmetron/dilaton models
 - 3) Vainshtein mechanism

massive gravity, Galileon models, braneworld models

Non-linearity of the Poisson equation for the fifth force leads to rich phenomenology