

A search for the origin of the interstellar comet 2I/Borisov

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ABSTRACT

The discovery of the second interstellar object 2I/Borisov on 2019 August 30 raises the question of whether it was ejected recently from a nearby stellar system. Here we compute the asymptotic incoming trajectory of 2I/Borisov, based on both recent and pre-discovery data extending back to December 2018, using a range of force models that account for cometary outgassing. From Gaia DR2 astrometry and radial velocities, we trace back in time the Galactic orbits of 7.4 million stars to look for close encounters with 2I/Borisov. The closest encounter we find took place 910 kyr ago with the M0V star Ross 573, at a separation of 0.068 pc (90% confidence interval of 0.053–0.091 pc) with a relative velocity of 23 km s⁻¹. This encounter is nine times closer than the closest past encounter identified for the first interstellar object 1I/‘Oumuamua. Ejection of 2I/Borisov via a three-body encounter in a binary or planetary system is possible, although such a large ejection velocity is unlikely to be obtained, and Ross 573 shows no signs of binarity. We also identify and discuss some other recent, close encounters, recognizing that if 2I/Borisov is more than about 10 Myr old, our search would be unlikely to find its parent system.

Key words. comets: individual (2I/Borisov) — comets: general

1. Introduction

The discovery of the first interstellar object, 1I/‘Oumuamua, on 2017 October 19 generated an intense period of observing using over 100 hours of medium to large telescope time for the two weeks following its discovery. The strong interest in 1I/‘Oumuamua resulted from this being the first opportunity to acquire detailed information from a fragment of material ejected from another star system. The scientific community was eager for the discovery of a second interstellar object because it would start to provide information on the diversity of interstellar objects.

The data obtained on 1I/‘Oumuamua immediately following its discovery focused on characterization of its physical properties, while the final data, obtained by the Very Large Telescope and Hubble Space Telescope between November 2017 and early January 2018, was an astrometric experiment designed to trace the orbit back to 1I/‘Oumuamua’s parent solar system. From these observations we learned that 1I/‘Oumuamua was small and red, typical of solar system comets, but looked asteroidal (without dust or gas, Meech et al. 2017; Ye et al. 2017; Fitzsimmons et al. 2018). The extreme brightness variations showed that the nucleus was rotating in an excited state and was highly elongated – something that is yet to be explained (Meech et al. 2017; Belton et al. 2018). Because of the lack of activity, it was surprising when the astrometry showed a strong non-gravitational signal that was perturbing the motion of 1I/‘Oumuamua as it left the solar system, suggesting that there was undetected cometary outgassing (Micheli et al. 2018). The non-Keplerian trajectory

was compared with the reconstructed orbits of 7.4 million stars from the Gaia DR2 catalog to try to identify past close encounters that might be 1I/‘Oumuamua’s parent star system, but no low velocity close encounters were found (Bailer-Jones et al. 2018a).

The enduring interest in interstellar objects has stimulated a rich field of investigation, with more than 120 papers written on 1I/‘Oumuamua to date. Yet because of its brief visit, we still lack information regarding some of the most fundamental questions about 1I/‘Oumuamua: What is it made of? How did it acquire its strange shape? Where did it come from?

Comet C/2019 Q4 (Borisov) was discovered at low elevation in the morning twilight on 2019 August 30 by Gennady Borisov at the MARGO observatory in Crimea (CBET 2019) when it was at 2.98 au from the Sun and had a short dust tail. After analyzing the astrometry reported to the Minor Planet Center (MPC) in the following week, the possibility of a highly hyperbolic orbit became apparent. Because it was moving more slowly than 1I/‘Oumuamua, it took longer for the orbit to be recognized as interstellar, but by 2019 September 10 our astrometric data from the Canada-France-Hawaii telescope showed a hyperbolic orbit at the 10- σ level with an eccentricity of 3.08 (later refined to 3.36). The hyperbolic orbit was published by the MPC on 2019 September 11 (Minor Planet Center 2019a) and this was officially designated as 2I/Borisov on 2019 September 24 (Minor Planet Center 2019b). 2I/Borisov is red (Jewitt & Luu 2019) with similar colors to 1I/‘Oumuamua; our data give $g-r = 0.64 \pm 0.01$ mag. Unlike 1I/‘Oumuamua, 2I/Borisov is actively outgassing, and CN has been detected, giving it a chemical composition that appears to be similar to solar system carbon-

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chain species depleted comets (Fitzsimmons et al. 2019; Opatom et al. 2019).

Pre-discovery observations subsequently reported by the Zwicky Transient Facility and going back as far as 2018 December 13 (Ye et al. 2019), brought the total observed arc to almost a year (we use here observations up to 2019 November 17) with a total of 916 observations.¹ In this paper we use these data to compute the asymptotic trajectory of 2I/Borisov, and use astrometry from the Gaia DR2 catalogue (Gaia Collaboration 2018) plus radial velocities to trace stellar positions back in time, in order to identify any close encounters. We cannot perform a survey that is “complete” in any useful sense of that term, because we are limited by the available Gaia astrometry and, more significantly, by the available matching radial velocities (details below). Our goal is therefore just to look for close encounters among specific stars, and not to attempt to correct for incompleteness in a statistical manner.

2. Asymptotic trajectory

To determine 2I/Borisov’s asymptotic incoming trajectory we fitted a suite of different models (listed below) to account for the different possible behavior of non-gravitational forces.

- Gravity-only (JPL solution 38).
- Marsden et al. (1973) model, in which the radial, transverse, and out-of-plane components of the non-gravitational perturbation are $A_i g(r)$ for $i = 1, 2, 3$ respectively. The A_i parameters are constant while the function $g(r)$, which describes the dependency on the heliocentric distance r , is driven here by H_2O . JPL solution 39 estimates only A_1 and A_2 , solution 40 also estimates A_3 , and solution 41 estimates all of these as well as a time offset relative to perihelion of the peak of the nongravitational acceleration (Yeomans & Chodas 1989).
- Marsden et al. (1973) model with $g(r)$ driven by CO (Meech & Svoren 2004). JPL solution 42 estimates only A_1 and A_2 , whereas solution 43 also estimates A_3 .
- Rotating jet model (Whipple 1950, 1951; Chesley & Yeomans 2005) with $g(r)$ driven by H_2O . The two jets are located at 150° and 135° of colatitude. There are two distinct minima for the pole: JPL solution 44 corresponds to a pole at $(\alpha, \delta) = (150^\circ, -55^\circ)$, solution 45 to $(\alpha, \delta) = (315^\circ, 25^\circ)$.
- Rotating jet model, with $g(r)$ driven by CO. JPL solution 46 corresponds to a pole at $(\alpha, \delta) = (205^\circ, -55^\circ)$, solution 47 to $(\alpha, \delta) = (340^\circ, 30^\circ)$ (Ye et al. 2019).

Similar to what Bailer-Jones et al. (2018a) did for II/Oumuamua, for each orbital solution we numerically integrated the trajectory back to 3000 BCE and then extrapolated to the Keplerian inbound asymptote relative to the solar system barycenter. The gravitational model takes into account the Sun, the eight planets, the Moon, Pluto, 16 main belt perturbors, and general relativity (Farnocchia et al. 2015). The parameters for these asymptotic solutions (giving the direction 2I/Borisov was coming from, and its speed, before encountering the Sun), including complete uncertainties, are listed in Table 1. A selection is plotted in Figure 1.

For the rest of our analysis we use only solution 47. The covariance for this solution in Table 1 corresponds to 1σ uncertainties in α , δ , and v_∞ of $3.5''$, $1.4''$, and 0.0049 km s^{-1} respectively. This solution captures reasonably well the overall scatter caused by the different non-gravitational models. The maximum

¹ https://www.minorplanetcenter.net/db_search/show_object?utf8=&object_id=2I

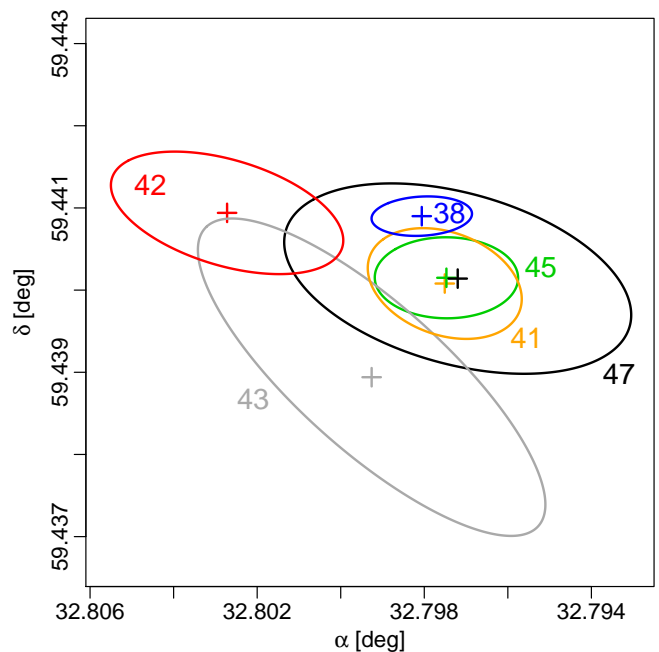


Fig. 1. Right ascension (α) and declination (δ) marked by a cross, plus uncertainty ellipse (90% confidence region), of the incoming asymptotic trajectory of 2I/Borisov. The labels refer to the JPL solution numbers in Table 1. Both axes span the same angular range of $\Delta\alpha \cos(\delta) = \Delta\delta = 0.00691^\circ$.

difference in the direction of the different asymptotic trajectories is $12''$, which corresponds to a transverse deviation of 0.002 pc over a path of 30 pc (the distance 2I/Borisov travelled to the closest encounter we find below). The maximum difference in the asymptotes’ velocity magnitudes is 8 m s^{-1} , corresponding to a displacement of 0.008 pc along the path after 1 Myr . Both of these are small compared to the uncertainties in the stellar trajectories, as we will see from the results. While revising this paper, further JPL solutions became available, yet these are all consistent with JPL solution 47 to well within the uncertainties, differing by less than 1 m s^{-1} and at most a few arcseconds. No further significant change of the incoming asymptote is expected.

3. Close stellar encounters

To identify close encounters, we integrate the orbits of a sample of stars back in time through a Galactic potential, and we do the same for 2I/Borisov starting from its asymptotic trajectory computed in the previous section. Our sample comprises all sources in Gaia DR2 that have 5-parameter astrometry, and radial velocities either from Gaia DR2, or from Simbad (Wenger et al. 2000) if not in Gaia DR2. This is mostly limited to cool stars (roughly $3500\text{--}7000 \text{ K}$) with $5 < G < 14 \text{ mag}$. We further limit our search to objects that have $\text{visibility_periods_used}^2 \geq 8$ and unit weight error (uwe) less than 35 (see section 2.2 of Bailer-Jones et al. 2018b for further explanation). In total our sample contains $7\,428\,838$ sources.³

² This is the number of distinct observation epochs, defined as an observation group separated from other groups by a gap of at least four days.

³ This total consists of three groups: $7\,039\,430$ sources have all six parameters from Gaia DR2; $337\,767$ sources have no radial velocity in

Table 1. Solutions for the incoming asymptotic trajectory (velocity vector) of 2I/Borisov. (α, δ) and v_∞ give the barycentric ICRF (International Celestial Reference Frame) direction and asymptotic velocity of 2I/Borisov. $\sigma(v_x)$, $\sigma(v_y)$, $\sigma(v_z)$ are the standard deviations of the Cartesian ICRF velocity components; the last three columns give the correlation coefficients ρ . For orientation, the Galactic coordinates of these solutions are $l = 132.9^\circ$, $b = -1.9^\circ$. We use just solution 47 in the analysis.

Solution	α [deg]	δ [deg]	v_∞ [km s ⁻¹]	$\sigma(v_x)$ [km s ⁻¹]	$\sigma(v_y)$ [km s ⁻¹]	$\sigma(v_z)$ [km s ⁻¹]	$\rho(v_x, v_y)$	$\rho(v_x, v_z)$	$\rho(v_y, v_z)$
38	32.79806	59.44090	32.284055	0.0001812	0.0002747	0.0004150	0.8210	0.8735	0.9878
39	32.79823	59.44022	32.284287	0.0003626	0.0003430	0.0005600	0.8465	0.8937	0.9761
40	32.79809	59.44016	32.284313	0.0003774	0.0003438	0.0005601	0.7925	0.8632	0.9725
41	32.79751	59.44008	32.283096	0.0005672	0.0005446	0.0011744	0.8654	0.9090	0.9664
42	32.80272	59.44094	32.279114	0.0006349	0.0003488	0.0011396	0.3297	0.8514	0.6946
43	32.79926	59.43894	32.285134	0.0017506	0.0006598	0.0022294	0.8543	0.9588	0.9173
44	32.79747	59.44015	32.282620	0.0005058	0.0005272	0.0010957	0.8971	0.9430	0.9636
45	32.79747	59.44015	32.282644	0.0004957	0.0005208	0.0010804	0.8989	0.9437	0.9646
46	32.79798	59.44021	32.284561	0.0018604	0.0015284	0.0041306	0.9886	0.9912	0.9925
47	32.79720	59.44014	32.286894	0.0024122	0.0011258	0.0041225	0.9282	0.9886	0.9500

As done in the 1I/Oumuamua study (Bailer-Jones et al. 2018a), we use a parallax zeropoint of -0.029 mas, i.e. this is subtracted from all parallaxes. The true parallax zeropoint may vary and may be slightly more negative for stars brighter than $G \simeq 16$ mag (e.g. Zinn et al. 2019), perhaps about -0.050 mas. This difference is generally smaller than the parallax uncertainties for our stars. Moreover, the uncertainties in the radial velocities tend to dominate the uncertainties in our inferred encounter parameters, so the exact choice of zeropoint is not critical.

We use the same Galactic model for the orbital integration as we did in the 1I/Oumuamua study. This is a smooth, three-component axisymmetric Galaxy model. Further details can be found in Bailer-Jones (2015). We do not include discrete components such as molecular clouds, not least because these evolve in unknown ways on the timescales of our integration (Myr), so cannot be correctly represented. Yet because of the limited distance horizon of our data, we generally trace orbits only for a few Myr (a few tens of pc), so the trajectories are virtually linear (demonstrated below). The exact choice of Galactic potential is therefore not very important. The accuracy of the resulting encounters is limited mostly by the stellar data.

Our method of finding encounters is identical to the one we used in Bailer-Jones et al. (2018a) to find possible parent stars of 1I/Oumuamua, where more details can be found. As in that work, in order to quantify the uncertainties in the encounter parameters, we generate 2000 surrogates for each star from its 6D covariance matrix (3D position, 3D velocity), and likewise for 2I/Borisov, and integrate all of these back in time.

We find 14 stars that approached within 1 pc of 2I/Borisov in the past. Table 2 shows the encounter parameters and other data on the stars. All of these have u_{we} less than a few, and so are formally good astrometric solutions (if they are not wide binary systems; see section 3.2). We are interested in encounters that are both very close, and very slow. As shown in Bailer-Jones et al. (2018a), a comet or asteroid ejected from a stellar system in a three-body interaction would have a characteristic ejection velocity of a few km s⁻¹ if the third body (the stellar companion) is a giant planet. We see in Table 2 much larger velocities, of

20 km s⁻¹ or more, which is only generally achievable with a stellar-mass companion.

3.1. Ross 573

The closest encounter we find passes very close, at a separation of just 0.068 pc (around 14 000 au) with a 90% confidence interval (CI) of 0.053–0.091 pc. The encounter took place 910 kyr ago at a relative speed of 23 km s⁻¹. Figure 2 shows the distribution of the surrogates. This object, Gaia DR2 5162123155863791744, has various other names including Ross 573, and is listed in Simbad as being an M0 dwarf with an apparent V-band magnitude of 10.0 mag. These agree with its Gaia DR2 apparent and absolute magnitudes and color (listed in Table 2), and effective temperature (T_{eff}) of 4050 K (68% CI 3980–4180 K). This is a high proper motion star (310 mas yr⁻¹) that is currently 21 pc from the Sun. None of the Gaia DR2 quality metrics suggest the Gaia astrometry is problematic (e.g. $u_{we} = 1.3$). No nearby companions are listed in Simbad or are apparent in images in Aladin (Bonnarel et al. 2000), so there is no evidence that the astrometric solution is compromised by duplicity. There are seven good HARPS spectra in the ESO archive taken over a period of nine years. Their radial velocities are consistent within 20 m s⁻¹, which is further evidence for a lack of significant binarity. The HARPS radial velocity is 14.1105 ± 0.0068 km s⁻¹. If we use this in the orbital integration, the encounter parameters hardly change: $t_{\text{enc}}^{\text{med}}$ decreases by 0.0039 pc and $t_{\text{enc}}^{\text{med}}$ increases by 41 kyr.

Assuming Ross 573 has an Oort cloud, and that the radius of such an Oort cloud is limited by the Galactic tide, then the maximum radius scales with the size of the star’s Hill sphere in the Galactic potential (Hanse et al. 2018). As Ross 573 is very close to the Sun, the scaling factor is simply $(M_{\text{star}}/M_\odot)^{1/3}$. With $M_{\text{star}} = 0.7 M_\odot$ from the spectral type, this scaling factor is 0.9. Thus the maximum radius for Ross 573’s Oort cloud is similar to that of the Sun’s, about 10^5 au. The encounter with 2I/Borisov is much less than this, so from this point of view Ross 573 is a plausible origin for this exo-comet.

Using the median encounter parameters and the adopted mass of Ross 573, we compute that 2I/Borisov was deflected by 36'' by this encounter. The typical relative velocities of, and timescales to, other close encounters in Table 2 are 20 km s⁻¹ and 500 kyr respectively, during which 2I/Borisov will have travelled 10 pc. Over this distance the deflection corresponds to a lateral

Gaia DR2 but do have a radial velocity and a Gaia DR2 source_id in Simbad (i.e. were cross matched by CDS); 51 641 sources have a radial velocity in Simbad but no Gaia DR2 source_id in Simbad, yet nonetheless have a close ($< 1''$) counterpart in Gaia DR2. This last group includes some duplicates of sources in the first group, but uses a Simbad as opposed to a Gaia DR2 radial velocity.

Table 2. The 14 stars that encounter 2I/Borisov in the past with a median encounter distance ($d_{\text{enc}}^{\text{med}}$) below 1 pc, sorted by this value. Column 1 reports a common name (Ross, Wolf, GJ, HD, HIP, in that order) from Simbad, if available (Gaia DR2 4828141619844398464 has no entry in Simbad). Columns 3, 6, and 9 are the median encounter time $t_{\text{enc}}^{\text{med}}$, median encounter distance $d_{\text{enc}}^{\text{med}}$, and median encounter velocity $v_{\text{enc}}^{\text{med}}$, respectively. The columns labeled 5% and 95% are the bounds of the corresponding 90% confidence intervals. Columns 12–17 list the parallax (ϖ); corrected for the zero-point), total proper motion (μ), and radial velocity (v_r) along with their 1- σ uncertainties. Columns 18, 19, and 20 are the apparent and absolute Gaia G-band magnitude (assuming zero extinction) and Gaia color respectively. As discussed in section 3.2, GJ 4384 is probably in a long-period binary orbit, so the encounter parameters shown here (which neglect this) are incorrect. [#] = HD 224635? (see text), [†] = G 7-34, [‡] = van Maanen's star, [§] = V* EV Lac.

name	Gaia DR2 source ID		t_{enc} [kyr]		d_{enc} [pc]		v_{enc} [km s ⁻¹]		ϖ mas	$\sigma(\varpi)$ mas	μ mas yr ⁻¹	$\sigma(\mu)$ mas yr ⁻¹	v_r km s ⁻¹	$\sigma(v_r)$ km s ⁻¹	G mag	M_G mag	BP-RP mag		
	med	95%	med	95%	med	95%	med	95%											
Ross 573	5162123155863791744	-909	-937	-882	0.068	0.053	0.091	22.6	22.0	23.3	47.57	0.03	309.22	0.09	13.1	0.4	9.33	7.72	1.80
GJ 4384 [#]	2875096978193873024	-1523	-1550	-1496	0.250	0.218	0.286	19.1	18.7	19.4	33.74	0.04	126.27	0.06	-7.9	0.2	6.32	3.96	0.69
GJ 3270 [†]	3299381442858615936	-444	-559	-365	0.459	0.377	0.580	32.1	25.5	39.0	68.58	0.08	400.33	0.17	14.0	4.0	12.25	11.43	3.13
HD 44867	3368960531532018816	-1424	-1450	-1396	0.572	0.360	1.128	85.9	85.6	86.2	8.00	0.09	48.61	0.11	71.2	0.2	6.04	0.56	1.20
	3338543951096093696	-2876	-2930	-2827	0.588	0.162	1.244	36.9	36.3	37.5	9.21	0.04	55.30	0.08	21.4	0.4	10.46	5.28	0.99
Wolf 28 [‡]	2552928187080872832	-14.7	-15.1	-14.3	0.647	0.629	0.668	284.0	275.4	292.3	231.77	0.04	2978.29	0.10	263.0	4.9	12.31	14.14	0.61
	4828141619844398464	-2401	-2507	-2301	0.675	0.559	0.824	64.5	61.7	67.2	6.35	0.02	39.52	0.04	77.2	1.7	12.99	7.00	1.37
GJ 1103A	3085716990368639744	-95	-105	-86	0.698	0.634	0.778	95.7	86.0	105.0	107.80	0.09	801.48	0.17	93.8	5.7	11.66	11.82	3.12
GJ 2070	3073508528645520000	-402	-579	-307	0.809	0.615	1.172	32.5	22.5	42.6	74.67	0.07	479.02	0.12	34.9	6.2	11.51	10.87	2.60
GJ 873 [§]	1934263333784036736	-177	-178	-176	0.818	0.813	0.823	27.6	27.4	27.7	198.04	0.04	842.10	0.09	0.4	0.1	9.00	10.48	2.73
	6743117625594496512	-1750	-1783	-1718	0.835	0.496	1.220	69.8	68.6	70.9	8.00	0.04	43.84	0.10	89.6	0.7	10.41	4.93	0.95
HD 27756	207166446152692736	-1283	-1297	-1269	0.859	0.799	0.924	35.3	35.0	35.7	21.58	0.06	81.70	0.10	8.3	0.2	9.05	5.72	0.99
HIP 34498	978301126629450368	-5283	-5354	-5211	0.944	0.487	1.770	25.1	24.8	25.3	7.40	0.04	33.68	0.10	1.0	0.1	8.97	3.32	0.96
HD 22781	217334764042444288	-889	-900	-878	0.969	0.900	1.034	35.8	35.4	36.2	30.67	0.11	102.62	0.18	8.3	0.2	8.52	5.96	1.08

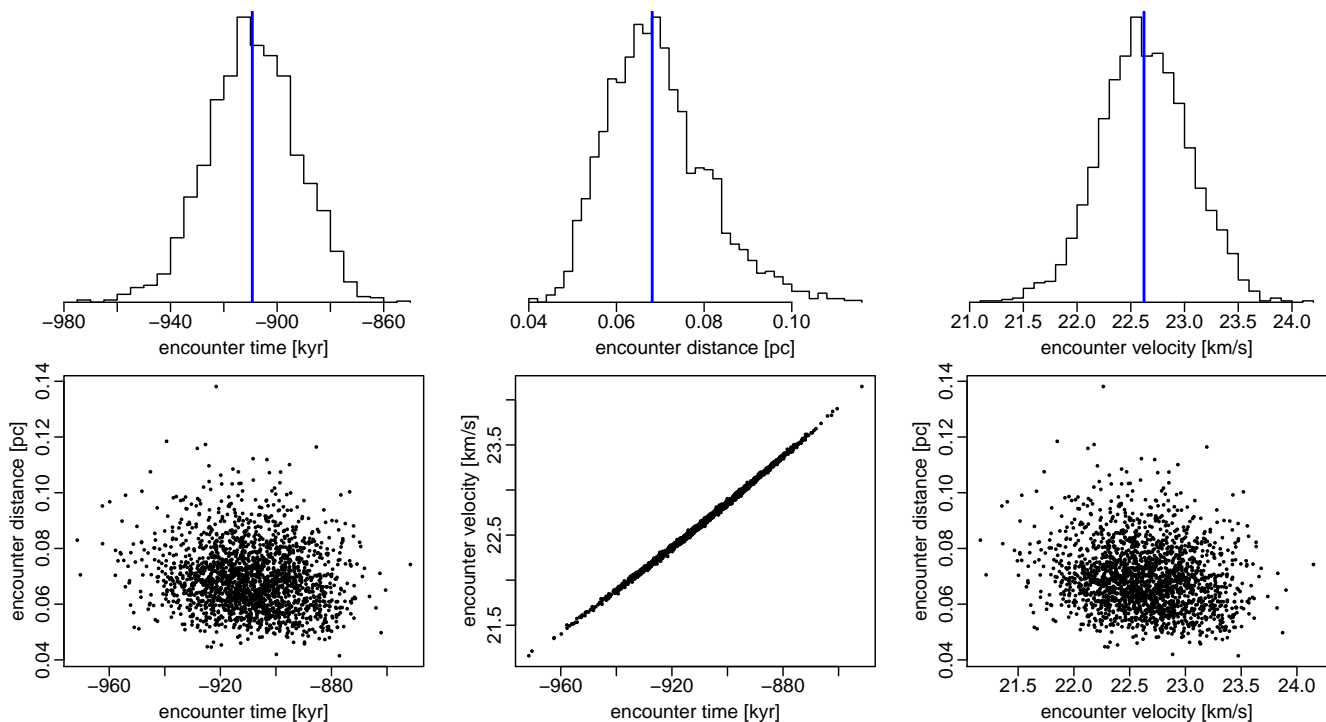


Fig. 2. Distribution over encounter parameters (for the 2000 surrogates) of 2I/Borisov asymptote solution 47 with Gaia DR2 5162123155863791744 (= Ross 573). The vertical blue lines in the upper panels show the medians of the distributions.

displacement of 0.002 pc. That is, our computation of other close encounters could be wrong by up to this amount due to our neglect of two-body effects. This is much smaller than the precision of the inferred encounter separations (and the other encounters induce even smaller errors).

The trajectory of Ross 573, and most of the other close encounters in Table 2, is so short that the exact Galactic potential has little impact on the encounter parameters. As an extreme test we set the potential to zero (i.e. use linear motions) and recompute the encounter parameters, again using all the surrogates. The median encounter time, distance, and velocity then change by -0.9 kyr, -0.0036 pc, and -0.03 km s $^{-1}$, respectively. These are much smaller than the uncertainties.

Ross 573 was also identified as a close encounter by Hallatt & Wiegert (2019) who infer a similar time and speed, but with a considerably larger encounter distance of 0.64 pc (90% CI of 0.60–0.69 pc). Their study used a combination of sources for the data on the stars, and it is not stated exactly which data they used for each object, although it seems likely that they used Gaia DR2 here. They use a slightly different Galaxy model, but given that our encounter parameters are hardly changed when we just assume linear motions, this is unlikely to be the reason for the discrepancy. More likely is the difference in the asymptotic trajectory of 2I/Borisov, in particular the velocity amplitude (see section 2). Hallatt & Wiegert (2019) used an early gravity-only solution obtained from JPL on 2 October 2019. Ours is based on a much longer data arc and accounts for the action of non-gravitational forces, so should be more accurate.

3.2. GJ 4384

Our second closest encounter is GJ 4384. To be more precise, the result in Table 2 uses astrometry for Gaia DR2 2875096978193873024 and the Simbad-listed radial velocity of -7.9 ± 0.2 km s $^{-1}$ (there is no Gaia DR2 radial velocity), as these objects match to within 0.1". Simbad lists a second source, HD 224635, as being only 0.01" from GJ 4384. However, from the data on these sources from various catalogues and publications listed on Simbad, we conclude that these two identifiers actually refer to the same, single source. Malkov et al. (2012), for example, in a study of binary systems, does not flag this as being either an eclipsing or spectroscopic binary star. It is also not identified as an unresolved binary in Hipparcos (Perryman & ESA 1997). We equate this star with ADS 17149A, HIP 118281, and WDS J23595+3343A, and refer to it here as s024 for short. Note that if we instead use the radial velocity listed in Simbad for HD 224635, which is -7.7 ± 2.0 km s $^{-1}$, then the encounter parameters shown in Table 2 are hardly changed ($d_{\text{enc}}^{\text{med}}$ by 0.0008 pc and $t_{\text{enc}}^{\text{med}}$ by 17 kyr).

GJ 4384 does, however, have a companion 2.475" away, Gaia DR2 2875096978193654528 (s528 for short), which also goes by the names HD 224636, ADS 17149B, and WDS J23595+3343B.

These two stars – s024 and s528 – both have good astrometric quality indicators, so their proximity should not have corrupted their astrometry (in agreement with expectations about Gaia for this separation). But they do have very similar parallaxes (the average is 33.82 mas and the difference 0.04 mas) and radial velocities (Simbad gives -4.6 ± 2.0 km s $^{-1}$ for HD 224636), so are almost certainly a physical binary. If they were in a common plane perpendicular to the line-of-sight, their physical separation would be 73 au and their relative velocity 4.8 km s $^{-1}$.

This compares to a relative orbital velocity of two solar-mass stars in a mutual circular orbit of this radius of 4.9 km s^{-1} (and their period would be 440 yr). As their proper motion difference of 35 mas yr^{-1} is about a quarter of their absolute proper motions, the velocity of their center-of-mass (c.o.m) could differ significantly from that of either component.

If we knew the masses of both components of the binary, we could solve for both the c.o.m and the Keplerian orbital elements (6+6 parameters from the 3D position and 3D velocity for each star). The spectral type of this pair is given by Struve & Franklin (1955) as F8+G1, with F8 presumably referring to GJ 4384 because this is slightly brighter and bluer (by 0.3 mag and 0.05 mag respectively in Gaia DR2). This agrees with their T_{eff} and positions in the color-absolute magnitude diagram, as given by Gaia DR2. We adopt a mass of $1 M_{\odot}$ for each component to compute the position and velocity of the c.o.m. Integrating the c.o.m's motion back in time, we find the time, distance, and relative speed of its closest encounter with 2I/Borisov to be -1380 kyr, 3.64 pc, and 20.9 km s^{-1} respectively. This is considerably more distant than the encounter found when neglecting the binary motion (the second line of Table 2). We do not compute the covariance on these parameters due to an inadequately-characterized correlation in the Gaia astrometric uncertainties for sources close to one another. The largest uncertainty is the radial velocity of s528. If we change this by $\pm 2\sigma$, the encounter distance remains between 3.3 and 4.0 pc. We conclude that the binary motion renders the GJ 4384–HD 224636 system a much more distant encounter than shown in Table 2. Although a binary system is an attractive option for ejecting 2I/Borisov, this encounter distance is far larger than the size of the system.

We note that GJ 4384 is listed as the second closest encounter to 2I/Borisov found by Hallatt & Wiegert (2019), with similar encounter parameters to ours ($d_{\text{enc}}^{\text{med}} = 0.32 \text{ pc}$) when we neglect binarity.

3.3. Other candidates

We note that several of our other close encounters are binary systems. One of these, HD 22781, has a planetary companion (Díaz et al. 2012), but the encounter velocity of 36 km s^{-1} is rather high for this planet to be plausibly involved in ejecting 2I/Borisov.

It is worth examining in addition those encounters beyond 1 pc (not listed in Table 2) that are particularly slow. The slowest encounter we find within 5 pc is with Gaia DR2 4769698316220254592 (= HD 34327), at 3.9 km s^{-1} (90% CI $3.7\text{--}4.2 \text{ km s}^{-1}$) but at 4.8 pc, 14.7 Myr ago. The next slowest encounter is nominally Gaia DR2 2007876324466455424 (= HD 239960A) with $v_{\text{enc}}^{\text{med}} = 5.6 \text{ km s}^{-1}$ (and d_{enc} between 1 and 3.5 pc). However, because of the large relative uncertainty in its velocity relative to 2I/Borisov, mostly arising from the radial velocity uncertainty, this encounter solution is poorly constrained.

The closest encounter that Hallatt & Wiegert (2019) find to 2I/Borisov is with GJ 3270, at $d_{\text{enc}}^{\text{med}} = 0.21 \text{ pc}$ (90% CI $0.08\text{--}0.32 \text{ pc}$) and $v_{\text{enc}}^{\text{med}} = 33 \text{ km s}^{-1}$. Our solution puts this encounter further away, with a 90% probability of approaching between 0.377 and 0.580 pc (see Table 2), and only a 1% probability that it approaches closer than 0.35 pc. The third closest encounter from Hallatt & Wiegert (2019), at 0.32–0.55 pc (90% CI), is with 2MASS J03552337+1133437 (= Gaia DR2 3303349202364648320). Our solutions put this much further away, at 1.07 pc (90% CI $0.98\text{--}1.17 \text{ pc}$), so is not listed in our table.

4. Discussion

This proximity of Ross 573's encounter with 2I/Borisov does not, on its own, mean the former is the origin of the latter. But we can estimate how often a star would happen to pass this close to 2I/Borisov. Using Gaia DR2, Bailer-Jones et al. (2018b) computed that the rate at which stars approach within 1 pc of the Sun is 19.7 ± 2.2 per Myr. This is for all types of stars and is corrected for Gaia's incompleteness. Bailer-Jones (2018) showed that this rate scales quadratically with the encounter distance. To first order, this same rate should apply for any object in the solar neighborhood, including 2I/Borisov. Thus the rate of encounters of stars within 0.067 pc of 2I/Borisov is expected to be 0.088 ± 0.010 per Myr, which is one every 11.3 Myr. This is 12 times longer than the time back to 2I/Borisov's encounter with Ross 573, suggesting this was unlikely to be just a chance encounter. However, because we only consider Gaia DR2 sources in the present study, a more appropriate comparison is with the encounter rate just for stars in Gaia DR2, i.e. before correcting for Gaia's incompleteness. This is 0.0083 encounters per Myr within 0.067 pc (Bailer-Jones et al. 2018b), which corresponds to an expected time of 120 Myr lapsing before such a close encounter, thus making it even less likely this was just a chance encounter.

An alternative estimate of the probability of coincidence can be made using Poisson statistics. Assuming a space density of stars (of potential origin) of 0.1 per cubic parsec, then the average number of flybys within 0.067 pc of 2I/Borisov over the distance it has travelled from Ross 573 to the Sun (30 pc) is $\lambda = 0.044$. The probability of 2I/Borisov having experienced an unassociated flyby as close as this is therefore $1 - e^{-\lambda} = 0.04$.

In their study of 1I/Oumuamua, Bailer-Jones et al. (2018a) showed that a comet could be ejected from a stellar system containing a stellar or giant planet companion with velocities of a few km/s. For the solar system – i.e. a system dominated by a $1 M_{\odot}$ star and a $1 M_{\text{Jup}}$ planet at 5 au – they found that the maximum ejection velocity of the comet was 17.4 km s^{-1} , which is not enough for 2I/Borisov. A more massive companion can eject at higher velocities, up to many tens of km s^{-1} for masses up to the hydrogen-burning limit. Yet to achieve the required ejection velocity of 23 km s^{-1} from Ross 573, a close interaction would be required, which is a priori less likely. Furthermore, there is no indication from the Gaia DR2 astrometry, HARPS spectroscopy, or the literature that Ross 573 is a binary. An alternative ejection mechanism is the close passage of another (unbounded) star. We searched Gaia DR2 for this. The only star coming within 1 pc of Ross 573 in a time window around the 2I/Borisov–Ross 573 encounter is Gaia DR2 347051160157809920, with encounter parameters $t_{\text{enc}} = -1080 \text{ kyr}$, $d_{\text{enc}} = 0.72 \text{ pc}$, $v_{\text{enc}} = 41 \text{ km s}^{-1}$, which is neither particularly close nor slow.

Our investigation is based on the assumption that 2I/Borisov is not older than about 10 Myr. This is because 2I/Borisov travels 33 pc per Myr relative to the Sun, yet the Gaia DR2+Simbad radial velocity sample becomes severely incomplete after a few hundred parsec. Currently we have no age for 2I/Borisov. Although its relatively large speed with respect to the local standard of rest – its pre-encounter space velocity is $(U, V, W) = (33.1, -6.8, 8.3) \text{ km s}^{-1}$ in our Galaxy model – could suggest it is old, this is only a weak indication. Future Gaia data releases will not extend the reliable traceback horizon substantially, because of the relatively bright magnitude limit on its radial velocities. Even if we had a much deeper six-parameter stellar catalogue, there remain limitations on how far back we can trace orbits, due to inaccuracies in our knowledge of the Galactic potential and its

evolution (e.g. Bailer-Jones et al. 2018a; Zhang 2018; Hallatt & Wiegert 2019).

In this study we have asked whether there are close encounters to specific stars among the available data. We have not attempted to estimate the overall rate of close encounters to 2I/Borisov by correcting for this incompleteness (as done for the Sun, for example, by García-Sánchez et al. 2001; Bailer-Jones 2018; Bailer-Jones et al. 2018b). Nor have we corrected for binary motion: even though this could be done in a few cases, the vast majority of binaries are either unidentified as binaries or could not be corrected using the available data.

5. Conclusions

We have found a very close (14 000 au; 90% CI 11 000–18 600 au) encounter between 2I/Borisov and the M0V star Ross 573 that took place around 900 kyr ago. If 2I/Borisov was ejected from orbit around this star via a three-body interaction, and assuming it has undergone no other encounter since, it must have been ejected at around 23 km s^{-1} . This is rather fast for ejection by a massive companion, for which there is also no observational evidence. We have not found a good candidate for a second star passing close to Ross 573 that could potentially have ejected 2I/Borisov. Thus although the probability of such a close encounter over the distance 2I/Borisov has travelled from this star to the Sun is low (0.04), there is no obvious mechanism to eject 2I/Borisov from the host star. A second initially interesting close approach was with the F8V star GJ 4384, 1500 kyr ago at 0.25 pc and 19 km s^{-1} . However, when taking into account its orbital motion with its solar-mass companion at a projected separation of 73 au, the encounter separation to the center-of-mass of this system increases to over 3 pc.

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